JWST Master Class Workshop

MOS hands-on

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NIRSpec MOS hands-on

Multi-object spectroscopy of distant galaxies



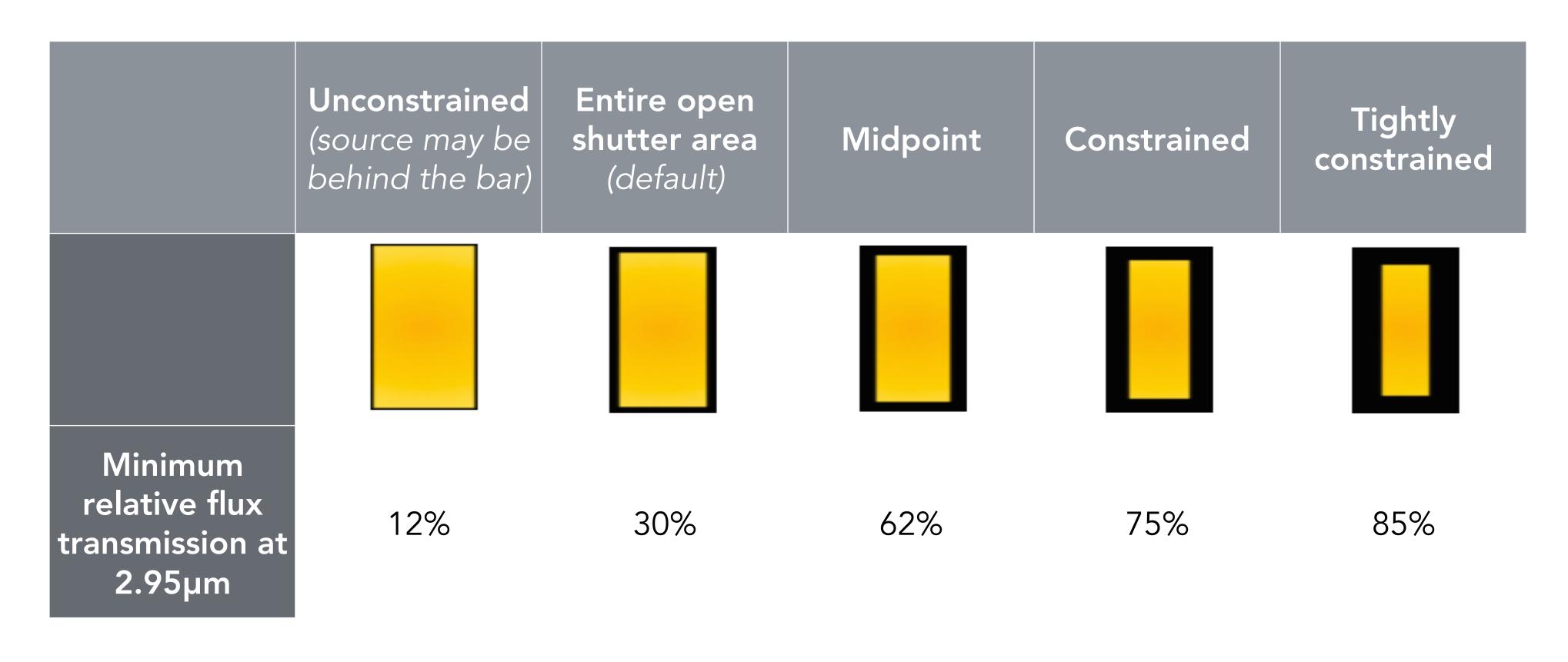
- In this hands-on session you will be asked to use
 - The Astronomer's Proposal Tool (APT)
 - More specifically, the MOS Planning Tool (MPT) which is part of the APT
- The same science case was already part of the ETC hands-on, so much of the observation planning is assumed to have been completed



Shutter planning constraints

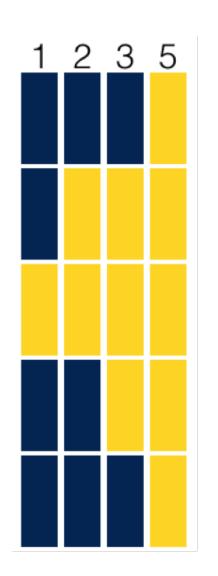
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- Slit loss is a function of the relative location in the shutter
- MPT allows adjustment of the constraint on where sources can be placed in shutters

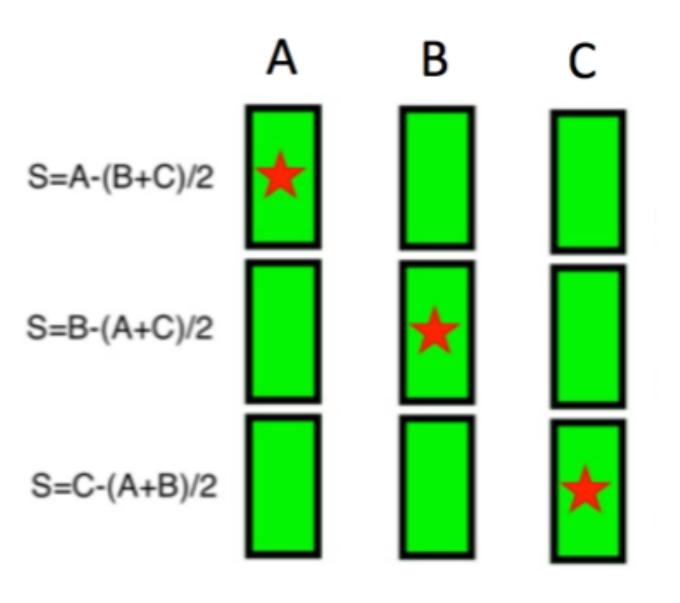


Nodding

- Open slitlets in cross-dispersion direction
- There are four selectable slitlets in MPT: 1, 2, 3, 5
- Telescope is repositioned slightly between exposures
 - No MSA re-configuration
 - Background subtraction between nodded exposures
 - Improves PSF sampling and bad pixel mitigation
- The spectrum from a slitlet will be segmented, with bar shadows between the individual shutters



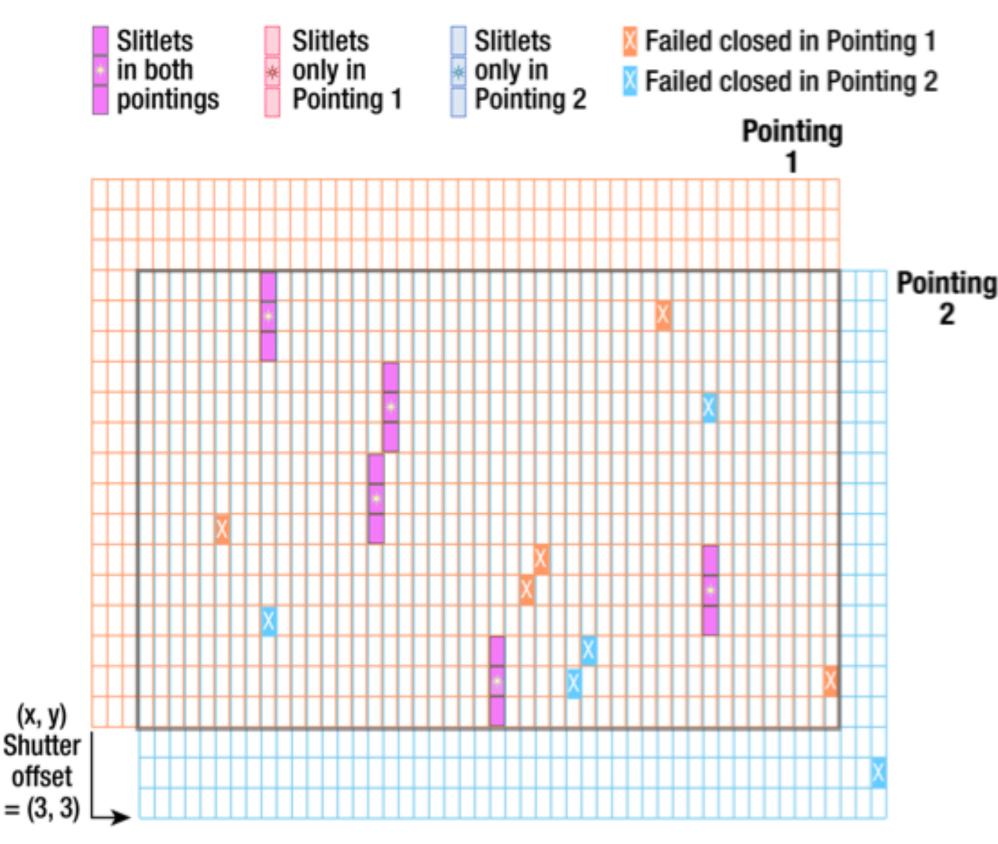




Dithering (fixed dither)

- Telescope repointed and MSA re-configured so many of the same sources fall in different shutters
 - Can be used to cover the wavelength gap
 - Improves PSF sampling and mitigates bad pixels
 - Allows observation of additional sources e.g. those behind shutter bars or the MSA mounting plate
 - Can mitigate effects of light leakage through MSA
- During optimisation, MPT will attempt to observe as many sources as possible at all dithers
- Nodding and dithering can be used alone or together

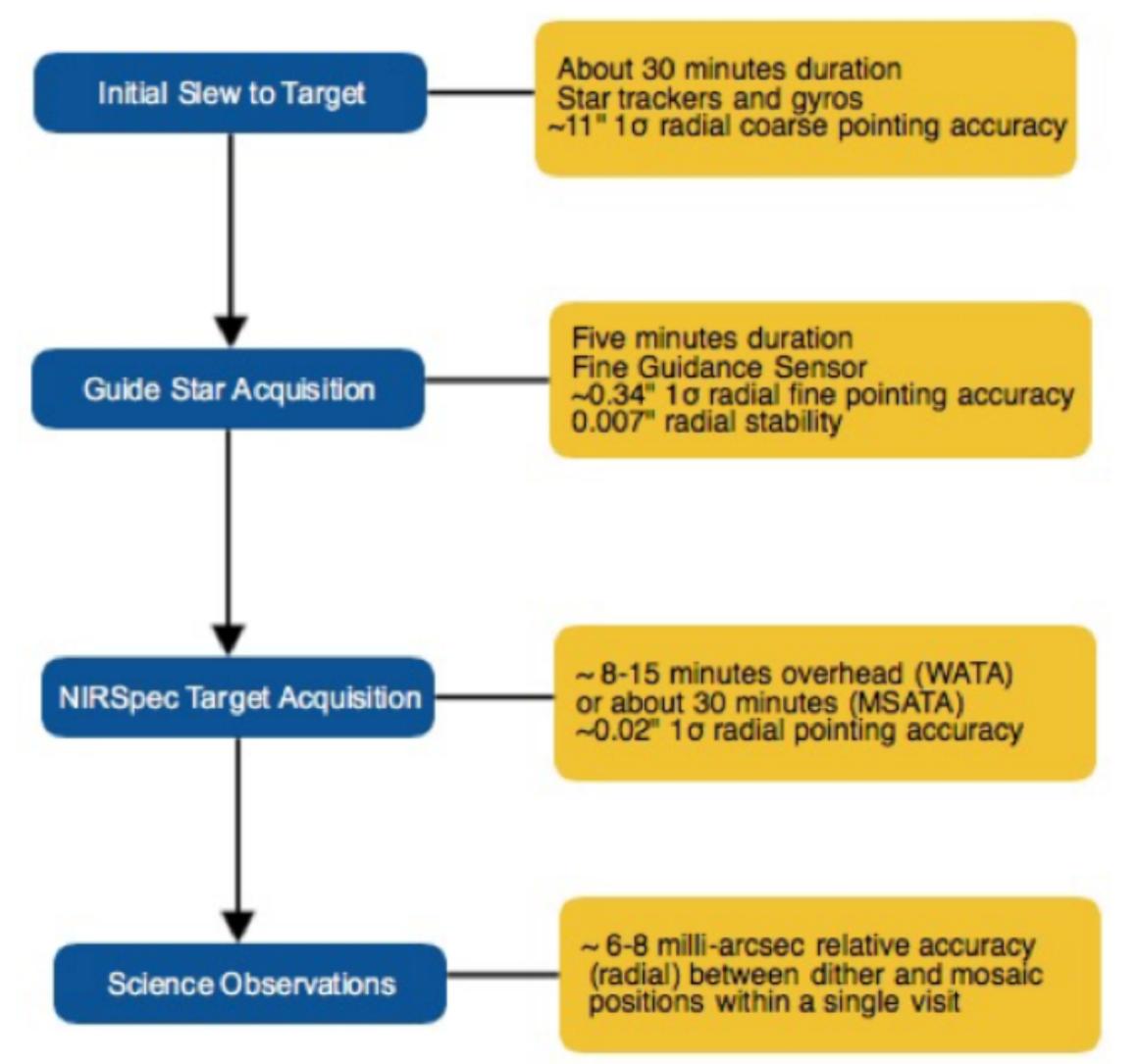




Target acquisition

ASTER CLASS MORRES

- MSATA is expected to be the usual TA method for the MOS
 - uses 5-8 reference stars over the entire MSA FoV
 - 8 chosen as a trade-off between TA accuracy and overheads
- MPT used to select reference stars at the assigned APA that will not be behind MSA bars or in failed shutters. This vetting is done at the Visit level at the first pointing in the Visit



Pre-imaging with NIRCam

- Is imaging available with
 - deep enough to identify sources?
 - wide enough to fill the MSA and plan reference stars?
 - accurate enough astrometry to plan MOS?
- If not, request NIRCam pre-imaging in your proposal
 - accurate enough astrometry to plan MOS?
 - NIRCam image should be large enough to allow for any NIRSpec APA
 - Ideally 5x5 arcmin
 - typically a 2x1 mosaic + dithers to cover gaps
 - NIRCam observations must be flight ready at proposal submission

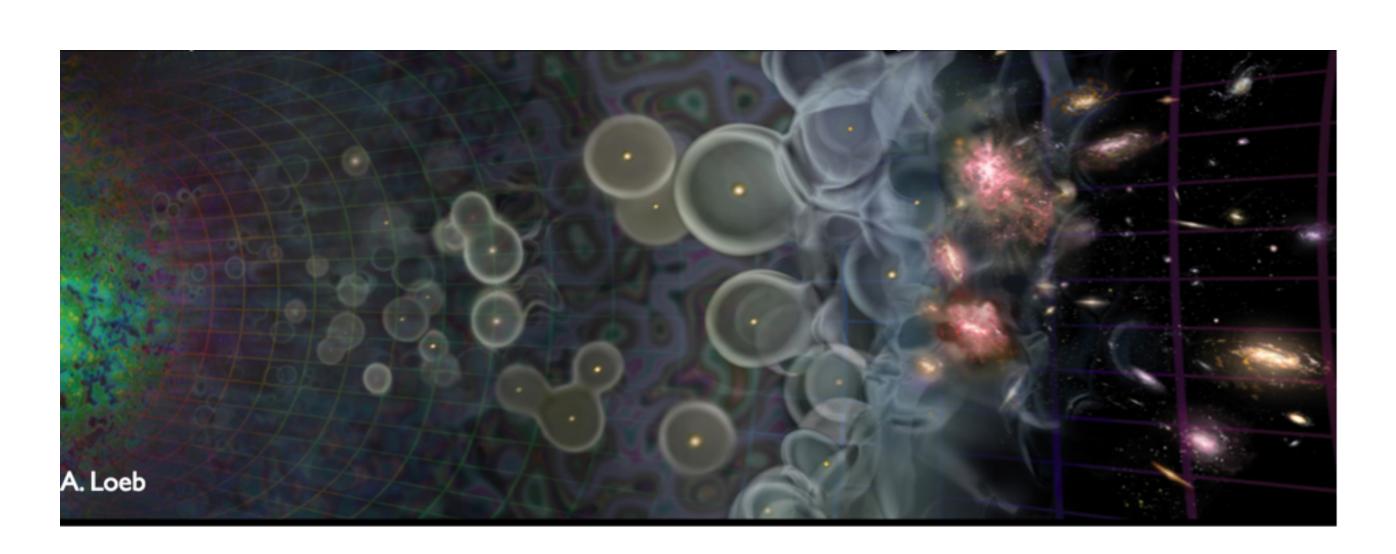




Science overview

Goal

- This program aims to study the evolution of galaxies from the early Universe (z>10), through the end of the dark ages (z=7-9), to the epoch of galaxy assembly (z=2-6)
 - Understanding very early stages of galaxy formation
 - Probing the epoch of re-ionization and the role of galaxies
 - Constraining the build-up of stellar mass and metals through time
 - Understanding the role of AGNs.
 - And looking for surprises...





Observation methodology

Methodology

- An in-depth program of this type would typically combine deep imaging (NIRCam) and follow-up spectroscopy (NIRSpec MOS)
- In the following, we will use a much simpler example:
 - deep NIRSpec MOS observations at a single location
 - input source catalog derived from existing HST imaging

Spectral configuration

- MOS at low and medium spectral resolution
- Type of sources
 - Galaxies over a wide range of redshifts handled as compact (point-like) objects
- Observation strategy
 - Combination of nodding (1x3 slitlets) and dithering



Instrument configuration

- Low spectral resolution (CLEAR/PRISM)
 - sensitivity to continuum
 - wavelength coverage (0.6-5.3 μm at once)
 - higher-multiplex MOS thanks to the shorter spectra on the detector
 - Main drawback: lack of spectral resolution...
- Medium spectral resolution (F100LP/G140M, F170LP/G235M, F290LP/G395M)
 - clean separation of emission lines
 - lace allows accurate centroid calculation for the lines
 - Main drawback(s):
 - 3 configurations required to cover 1.0-5.2 μm
 - high multiplex only possible if spectra allowed to overlap
- Configurations give complementary information: include all 4

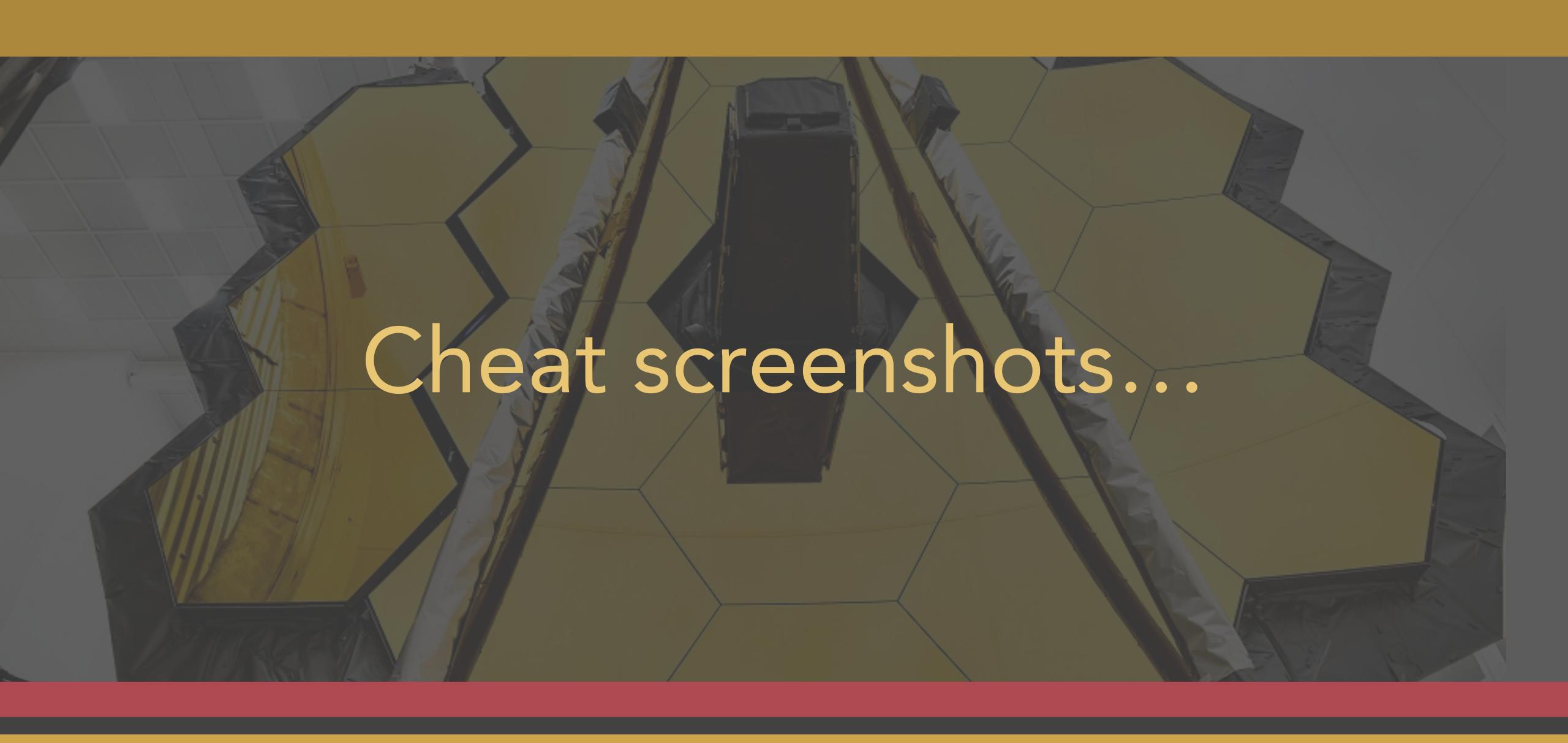




Getting started

- Follow the steps on the handout
 - Plan your observation strategy
 - Open APT
 - Load the catalogue
 - Set-up the MPT Planner
 - Generate the plan
 - Create the observation





Plan your observation

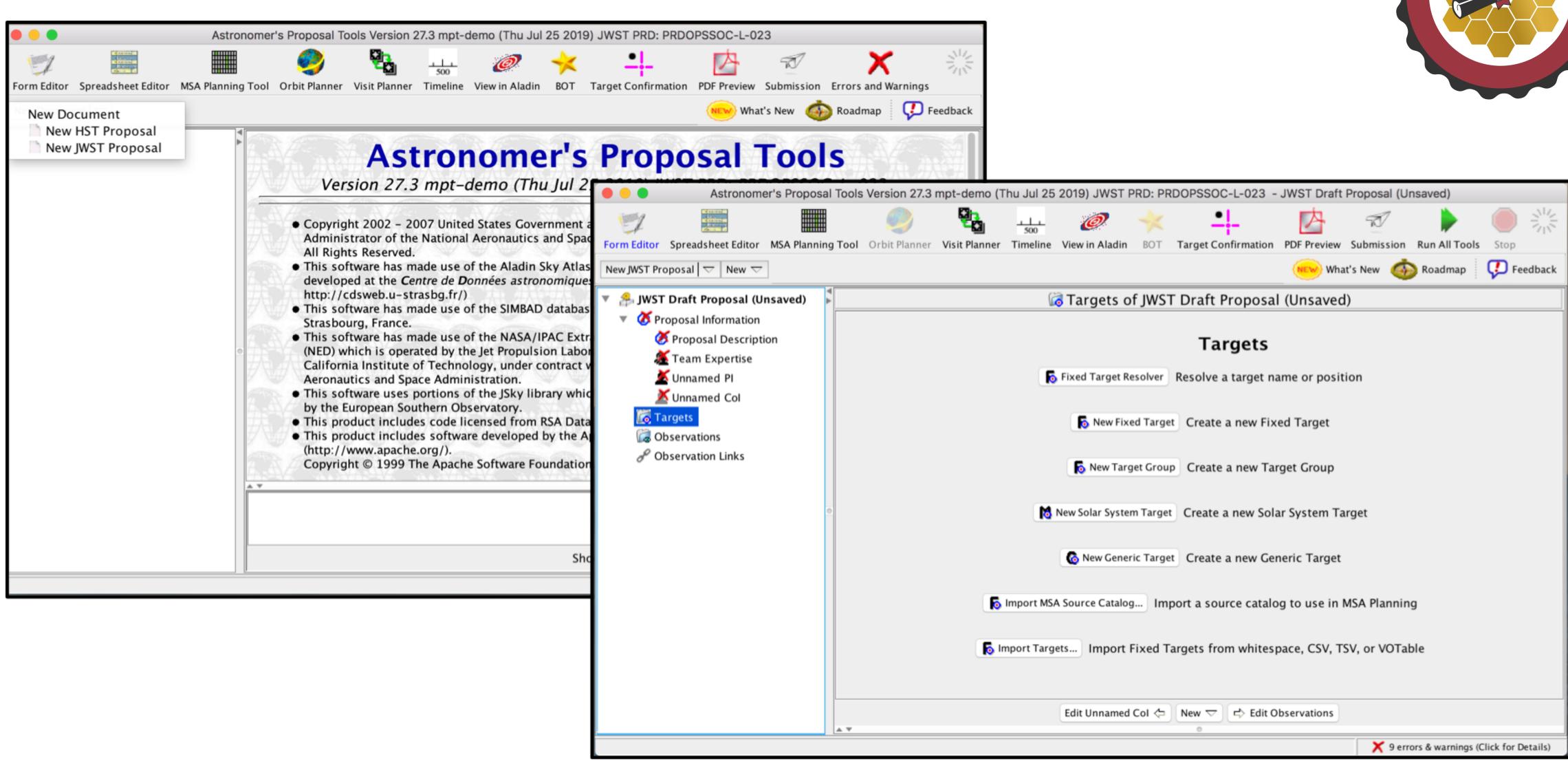


Dither	Nod	CLEAR/PRISM	F100LP/G140M	F170LP/G235M	F290LP/G395M
1	1	8	3	3	3
	2	8	3	3	3
	3	8	3	3	3
2	1	8	3	3	3
	2	8	3	3	3
	3	8	3	3	3
	1	8	2	2	2
3	2	8	2	2	2
	3	8	2	2	2
TOTAL Exps		72	24	24	24

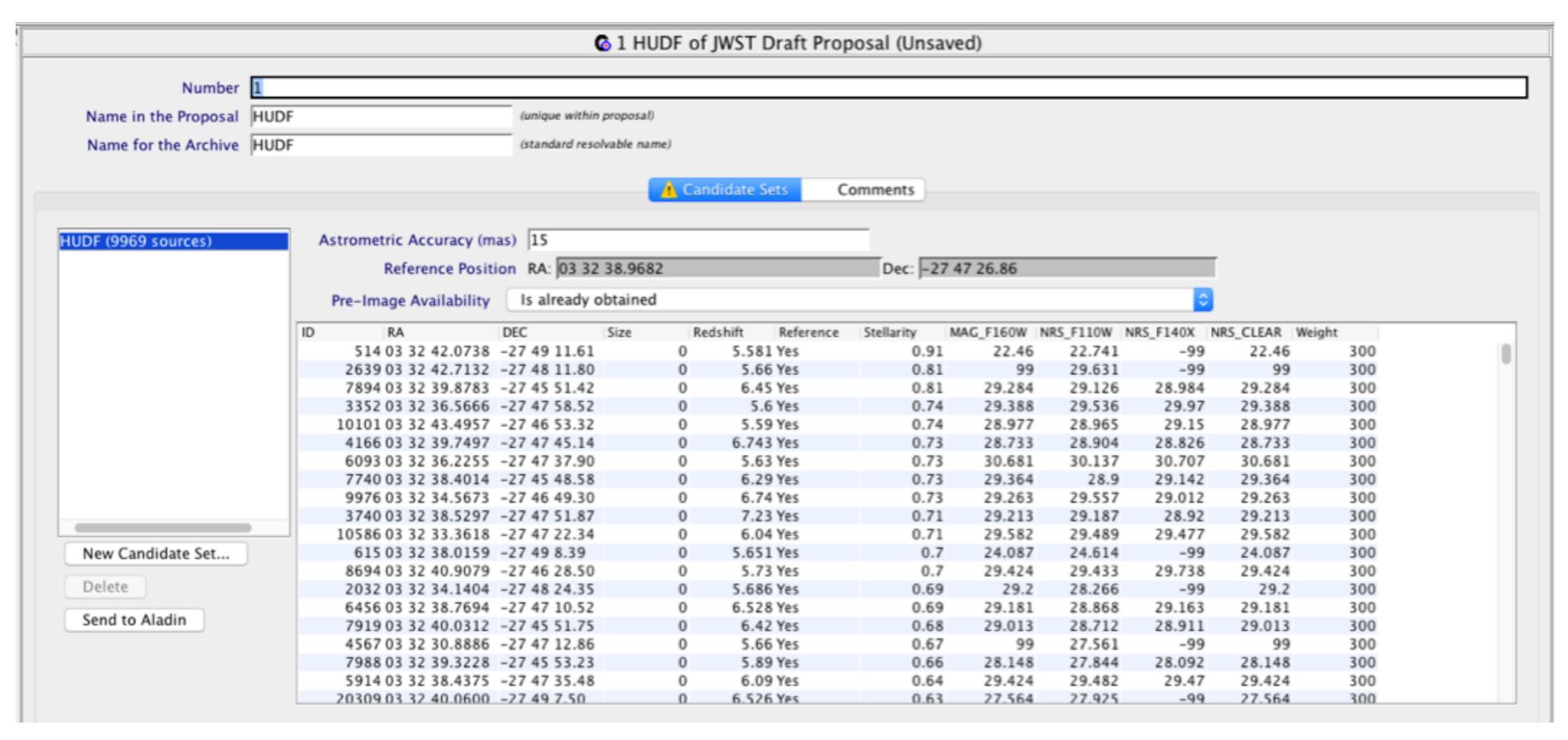
• NOTE: for simplicity, create a program with 27 exposures per medium-resolution grating

Open APT



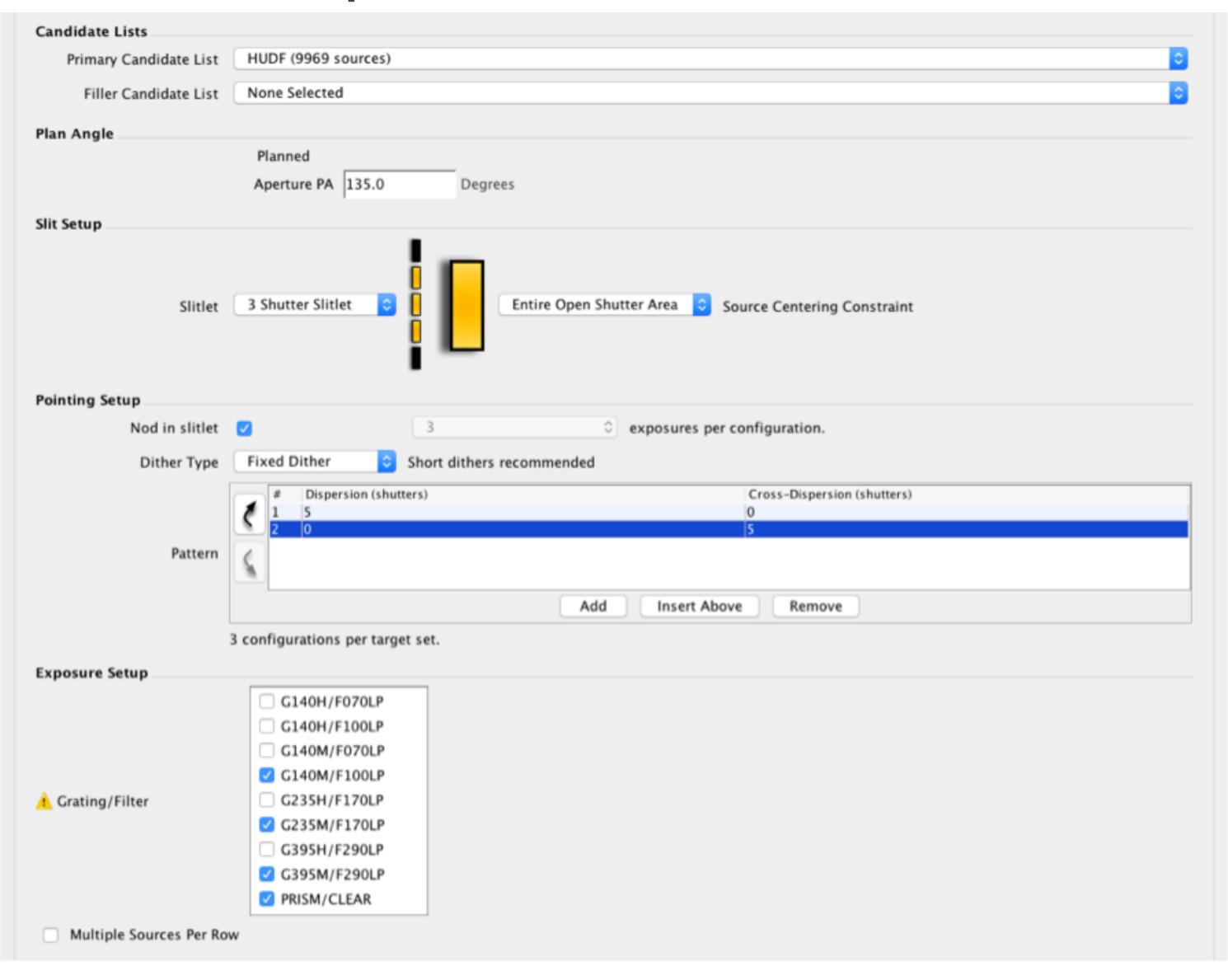


Load the catalogue





Set-up the MPT planner

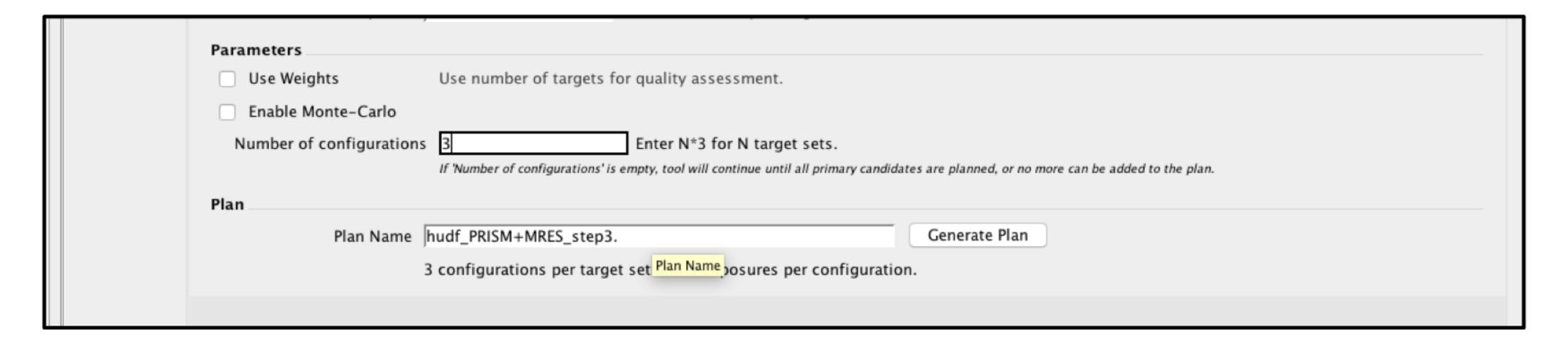




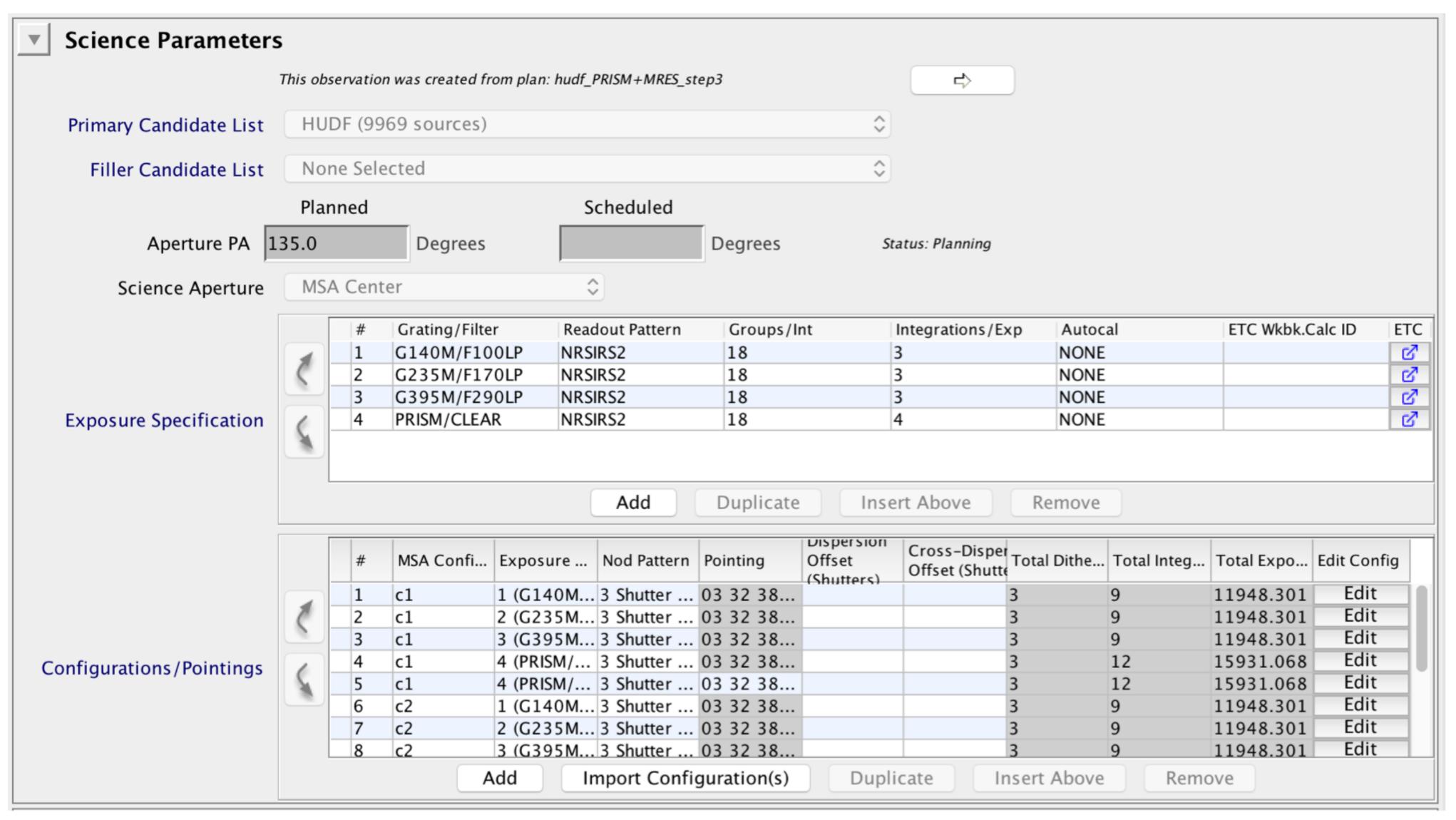
Planner parameters



Search Grid	Search Grid		
	Search Area Dimensions:		
Center	RA: 03 32 39.0067	Dec: -27 47 29.39	
Width	40	Arcseconds	
Height	40	Arcseconds	
Search Step Size	3.0	Arcseconds. 225 pointings will be tested.	



Create the observation







MOS/MPT Help and JDox

- JWST Help Desk: https://jwsthelp.stsci.edu
- JDox Home page: https://jwst-docs.stsci.edu/
 - Updated for the new version of APT/MPT
- Specific MOS/MPT materials:
 - ► MOS Roadmap
 - Multi-Object Spectroscopy
 - NIRSpec MOS Observing Process
 - NIRSpec MOS Operations
 - Catalogues and Images
 - Pre-Imaging Using NIRCam
 - Confirmation Images
 - Slit Losses
 - ► NIRSpec MOS Recommended Strategies
 - ► MOS APT Template

